

The ALMA-PILS survey: Detection of CH₃NCO toward the low-mass protostar IRAS 16293-2422 and laboratory constraints on its formation

N. F. W. Ligterink,^{1,2*} A. Coutens,^{3†} V. Kofman,¹ H. S. P. Müller,⁴ R. T. Garrod,⁵ H. Calcutt,⁶ S. F. Wampfler,⁷ J. K. Jørgensen,⁶ H. Linnartz,¹ E. F. van Dishoeck^{2,8}

¹*Sackler Laboratory for Astrophysics, Leiden Observatory, Leiden University, PO Box 9513, 2300 RA Leiden, The Netherlands*

²*Leiden Observatory, Leiden University, PO Box 9513, 2300 RA Leiden, The Netherlands*

³*Department of Physics and Astronomy, University College London, Gower St., London, WC1E 6BT, UK*

⁴*I. Physikalisches Institut, Universität zu Köln, Zùlpicher Str. 77, 50937 Köln, Germany*

⁵*Departments of Chemistry and Astronomy, University of Virginia, Charlottesville, VA 22904, USA*

⁶*Centre for Star and Planet Formation, Niels Bohr Institute & Natural History Museum of Denmark, University of Copenhagen, Øster Voldgade 5*

⁷*Center for Space and Habitability, University of Bern, Sidlerstrasse 5, CH-3012 Bern, Switzerland*

⁸*Max-Planck Institut für Extraterrestrische Physik (MPE), Giessenbachstr. 1, 85748 Garching, Germany*

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ABSTRACT

Methyl isocyanate (CH₃NCO) belongs to a select group of interstellar molecules considered to be relevant precursors in the formation of larger organic compounds, including those with peptide bonds. The molecule has only been detected in a couple of high-mass protostars and potentially on comets. A formation route on icy grains has been postulated for this molecule but experimental evidence is lacking. Here we extend the range of environments where methyl isocyanate is found, and unambiguously identify CH₃NCO through the detection of 43 unblended transitions in the ALMA Protostellar Interferometric Line Survey (PILS) of the low mass solar-type protostellar binary IRAS 16293-2422. The molecule is detected toward both components of the binary with a ratio HNC/CH₃NCO ~4–12. The isomers CH₃CNO and CH₃OCN are not identified, resulting in upper abundance ratios of CH₃NCO/CH₃CNO > 100 and CH₃NCO/CH₃OCN > 10. The resulting abundance ratios compare well with those found for related N-containing species toward high-mass protostars. To constrain its formation, a set of cryogenic UHV experiments is performed. VUV irradiation of CH₄:HNC mixtures at 20 K strongly indicate that methyl isocyanate can be formed in the solid-state through CH₃ and (H)NCO recombinations. Combined with gas-grain models that include this reaction, the solid-state route is found to be a plausible scenario to explain the methyl isocyanate abundances found in IRAS 16293-2422.

Key words: Astrochemistry - ISM: individual objects: IRAS 16293-2422 - ISM: molecules - Methods: laboratory: molecular - Techniques: spectroscopic - Molecular processes

1 INTRODUCTION

Complex organic molecules, defined in astrochemistry as molecules that consist of six or more atoms of which at least one is a carbon atom, are widely found in star-forming regions (Herbst & van Dishoeck 2009). A special category

of complex molecules is that of the prebiotics, molecules that can be linked via their chemical structures or reactivity to life bearing molecules, such as amino-acids and sugars. Methyl isocyanate, CH₃NCO, also known as isocyanomethane, is a molecule that falls in this category, because of its structural similarity with a peptide bond (Figure 1). This type of bond connects amino-acids to form proteins and as such is interesting because it connects to chemistry relevant to the formation of the building blocks of life.

* E-mail: ligterink@strw.leidenuniv.nl

† E-mail: a.coutens@ucl.ac.uk

Methyl isocyanate Isocyanic acid Generalized peptide bond

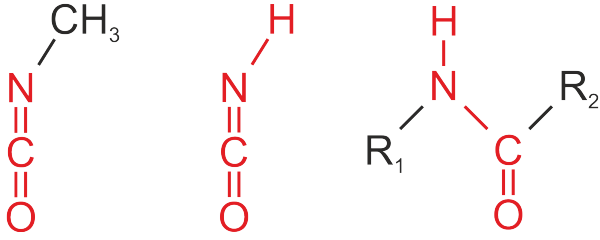


Figure 1. Structure of methyl isocyanate (left), isocyanic acid (middle) and the generalized structure of a peptide bond (right). In each structure the components of the peptide bond are highlighted in red. R_1 and R_2 are different molecular functional groups, which can, among other possibilities, be a methyl group as for methyl isocyanate.

The majority of identified complex molecules has mainly been detected toward high-mass hot cores, such as Orion KL and Sgr B2 (e.g., Blake et al. 1987, Nummelin et al. 2000, Belloche et al. 2013, 2014, Tercero et al. 2013, Neill et al. 2014, Crockett et al. 2014), but over the past decades detections toward low mass, sun-like, protostars such as IRAS 16293-2422 (hereafter IRAS16293) have been regularly reported as well. IRAS16293 ($d = 120$ pc) is considered as a protostellar template for low-mass sources and is particularly rich in organic molecules (van Dishoeck et al. 1995, Cazaux et al. 2003, Bottinelli et al. 2004, Kuan et al. 2004, Bisschop et al. 2008, Jaber et al. 2014). Jørgensen et al. (2012) demonstrated the capabilities of the Atacama Large Millimeter/submillimeter Array (ALMA) with the detection of the prebiotic molecule glycolaldehyde (see also Jørgensen et al. 2016, for a history of chemical studies of this source). More recently, other complex molecules (acetone, propanal, ethylene oxide) were identified in the framework of the PILS survey (Lykke et al. 2017). Even the deuterated isotopologues of several complex molecules were detected toward this source (Parise et al. 2003, Coutens et al. 2016, Jørgensen et al. 2016).

Unlike other molecules such as isocyanic acid (HNCO) and formamide (NH_2CHO) (Bisschop et al. 2007, López-Sepulcre et al. 2015, Coutens et al. 2016), which have a similar peptide-like structure as methyl isocyanate, CH_3NCO only recently started to attract attention in the astrochemistry community. This activity was inspired by a report from Goesmann et al. (2015) that CH_3NCO may be abundantly present on the surface of comet 67P/Churyumov-Gerasimenko, as measured with the *Cometary Sampler and Composition* (COSAC) instrument of Rosetta’s Philae lander. Its measured high abundance of 1.3% with respect to H_2O and $\text{CH}_3\text{NCO}/\text{HNCO}=4.3$ was based on the assumption that the mass peak recorded at $m/z = 57$ is dominated by CH_3NCO , as COSAC’s low mass resolution does not allow discrimination of different species with nearly identical mass. Recent measurements with the *Rosetta Orbiter Spectrometer for Ion and Neutral Analysis* (ROSINA) at much higher mass resolution showed that the CH_3NCO coma abundance is significantly lower (K. Altwegg, private communication). Nevertheless, the possibility of linking complex molecules in star- and planet-forming regions with those

present in comets triggered the search for methyl isocyanate in the interstellar medium.

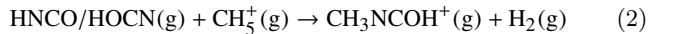
Interstellar CH_3NCO was first detected toward Sgr B2(N) by Halfen et al. (2015) at low rotational temperatures of ~ 25 K with a column density ratio of $N(\text{HNCO})/N(\text{CH}_3\text{NCO}) = 35\text{--}53$ depending on the specific velocity component. Cernicharo et al. (2016) detected methyl isocyanate toward Orion KL at $N(\text{HNCO})/N(\text{CH}_3\text{NCO}) \approx 15$ and $T_{\text{ex}} = 150$ K. Their observations toward the cold prestellar core B1-b did not yield a detection of the molecule down to an upper column density limit of $< 2 \times 10^{11} \text{ cm}^{-2}$ or $N(\text{HNCO})/N(\text{CH}_3\text{NCO}) > 42$ (based on a HNCO detection toward the same source by López-Sepulcre et al. 2015). In the same paper, publicly accessible Sgr B2 observations from Belloche et al. (2013) were re-analysed with an updated spectral line list and, interestingly, yielded a detection of warm methyl isocyanate at $T_{\text{ex}} \approx 200$ K and $N(\text{HNCO})/N(\text{CH}_3\text{NCO}) \approx 40$. Attempts to also identify the methyl isocyanate isomer CH_3CNO were unsuccessful down to $N(\text{CH}_3\text{NCO})/N(\text{CH}_3\text{CNO}) > 100$.

The astrochemical origin of methyl isocyanate is not yet understood and this is partly due to the limited number of laboratory studies that have been performed. Henderson & Gudipati (2015) tentatively assigned a mass fragmentation peak to CH_3NCO after VUV irradiating solid-state mixtures of $\text{NH}_3:\text{CH}_3\text{OH}$. In other experiments by Ruzi & Anderson (2012) UV irradiation of frozen *n*-methylformamide (CH_3NHCHO) also yielded methyl isocyanate, although it was concluded to represent a minor product channel.

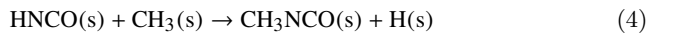
A number of formation routes have been hypothesized by astrochemists. Halfen et al. (2015) postulated gas-phase formation by HNCO or HOCN methylation:



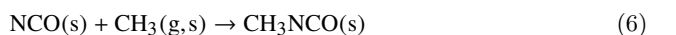
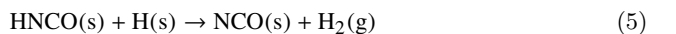
or reactions of HNCO or HOCN with protonated methane, followed by electron recombination:



Cernicharo et al. (2016) favoured solid-state formation mechanisms based on the detection of CH_3NCO toward hot cores and its non-detection in the cold dark cloud B1-b. Particularly, the methylation of HNCO has been mentioned as a possible route to form methyl isocyanate in the solid-state, i.e., on the surface of an icy dust grain:



Belloche et al. (2017) used the grain-surface radical-addition reaction $\text{CH}_3 + \text{NCO} \rightarrow \text{CH}_3\text{NCO}$ in their models, with most of the NCO formed via H-abstraction of HNCO:



These postulated routes require the reactants to be

present in sufficient amounts. Gaseous HNCO is detected in high abundances in protostellar environments and has been imaged in IRAS16293, showing it to be prominent in both source A and B (Bisschop et al. 2008, Coutens et al. 2016). It likely results from sublimation of OCN⁻, known to be a major ice component in low-mass protostellar envelopes (van Broekhuizen et al. 2005). A significant abundance of CH₃ gas is a more speculative assumption since the molecule can only be observed by infrared spectroscopy and has so far only been seen in diffuse gas toward the Galactic Center (Feuchtgruber et al. 2000). Alternatively, CH₃ radicals can be produced *in situ* in ices by photodissociation of known abundant ice components like CH₄ or CH₃OH and then react with HNCO or OCN⁻. This is the solid-state route that is investigated here.

In this work we present the first detection of methyl isocyanate toward both components of the low mass protobinary IRAS16293 on scales of < 100 AU. An independent detection toward source B is also reported by Martín-Doménech et al. (2017). A set of systematic laboratory experiments is presented in order to validate the solid-state formation routes of CH₃NCO. The observational work is presented in Section 2 and the laboratory work in Section 3. In Section 4 the results of the observations and laboratory experiments are compared and discussed in the context of recent astrochemical models. The conclusions of this paper are given in Section 5.

2 OBSERVATIONS

2.1 The ALMA PILS survey

We searched for methyl isocyanate in the Protostellar Interferometric Line Survey (PILS) data, an unbiased spectral survey of the low-mass protostellar binary IRAS16293 with ALMA. A full description and data reduction of the survey is presented in Jørgensen et al. (2016). Briefly, this survey covers a spectral range from 329.147 to 362.896 GHz and was obtained with both the 12m array and the Atacama Compact Array (ACA). The beam size ranges between $\sim 0.4''$ and $0.7''$ depending on the configuration at the time of the observations. The rms of the combined data sets is about 7–10 mJy beam⁻¹ channel⁻¹, i.e., approximately 4–5 mJy beam⁻¹ km s⁻¹. To facilitate the analysis, the combined data set used in this paper was produced with a circular restoring beam of $0.5''$ at a spectral resolution of 0.2 km s⁻¹.

Two positions are analysed in this study. The first position is offset by one beam diameter ($\sim 0.5''$) from the continuum peak of source B in the south west direction ($\alpha_{J2000}=16^{\text{h}}32^{\text{m}}22^{\text{s}}.58$, $\delta_{J2000}=-24^{\circ}28'32.8''$) (see high resolution images in Baryshev et al. 2015). Source B presents narrow lines (FWHM ~ 1 km s⁻¹). This position is found to be optimal for line identifications, as the lines are particularly bright, do not have strong absorption features toward the bright continuum of source B, and do not suffer from high continuum optical depth (Coutens et al. 2016, Lykke et al. 2017, Jørgensen et al. 2016). In this paper, we also analyse source A, which exhibits broader lines than source B making the line identification quite challenging (Pineda et al. 2012). The linewidth varies, however, depending on

the spatial separation from this source. With an average FWHM of ~ 2.5 km s⁻¹, the position $\alpha_{J2000}=16^{\text{h}}32^{\text{m}}22^{\text{s}}.90$, $\delta_{J2000}=-24^{\circ}28'36.2''$ appears to be one of the best positions to search for new species toward source A ($0.3''$ offset). At this position, the emission is centered at v_{LSR} of ~ 0.8 km s⁻¹, blueshifted from the source A velocity of $v_{\text{LSR}} = 3.2$ km s⁻¹ (Jørgensen et al. 2011).

Bisschop et al. (2008) found HNCO and other nitrogen containing species such as CH₃CN to be more prominent toward source A than source B. Consequently, it is also interesting to check whether there exists a small scale chemical differentiation among the N-bearing species for the two sources.

2.2 Results

Methyl isocyanate is detected toward both components A and B of IRAS16293. The identification is based on spectroscopic data from the Cologne Database for Molecular Spectroscopy (CDMS, Müller et al. 2001, 2005), taken from Cernicharo et al. (2016) as well as from Koput (1986). CH₃NCO is an asymmetric rotor with the NCO group lying at an angle of 140° and a low-lying (~ 182 cm⁻¹) C-N-C bending mode ν_b . Moreover, the barrier to internal rotation of the CH₃ group is low, only 21 cm⁻¹. The pure rotational spectrum therefore has *A* and *E* torsional states and vibrationally excited transitions can become detectable at temperatures of a few hundred K. The labelling of the states used here refers to the quantum numbers *J* (rotational angular momentum), *K_a* and *K_c* (projection of angular momentum on the respective inertial axes) with internal rotation interactions indicated by the quantum number *m*, with *m*=0 and ± 3 for the *A* states and *m*=1, -2 and 4 for the *E* states (Halfen et al. 2015, Cernicharo et al. 2016).

Using the CASSIS software¹, we have been able to identify 43 unblended lines of CH₃NCO in the bending ground state ($\nu_b = 0$) with upper energy levels E_{up} ranging from 320 to 670 K toward source B (see Table A1). Figure 3 shows the unblended lines detected toward this component as well as the LTE modeling for two different excitation temperatures at 100 and 300 K. Both excitation temperatures allow to reproduce the observations; only predicted transitions are observed. Our data are not sensitive to any cold CH₃NCO component since lines with low E_{up} values are missing in the spectral range of the PILS survey. Among the complex molecules that were detected and analysed towards source B, some (acetaldehyde, ethylene oxide) exhibit a relatively low excitation temperature of ~ 125 K (Lykke et al. 2017), while others (formamide, isocyanic acid, methanol, methyl formate, glycolaldehyde, ethylene glycol) show a higher excitation temperature of ~ 300 K (Jørgensen et al. 2012, Coutens et al. 2016, Jørgensen et al. 2016). Their spatial distribution is however rather similar and it is not possible to determine to which category CH₃NCO belongs. The methyl isocyanate column density is not very sensitive to the exact value of the excitation temperature: assuming the same source size of $0.5''$ as used in previous PILS studies (Coutens et al. 2016, Jørgensen et al. 2016, Lykke et al. 2017), the CH₃NCO column density is found to be about

¹ <http://cassis.irap.omp.eu/>

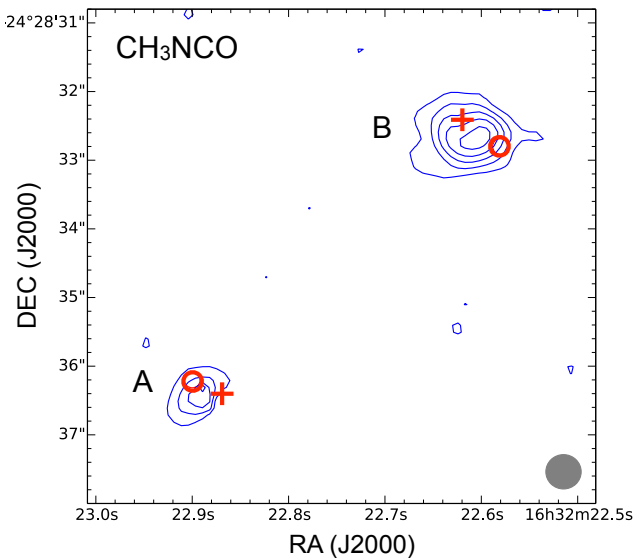


Figure 2. Integrated intensity map of the CH_3NCO $39_{0,39,0}-38_{0,38,0}$ transition at 336339.9 MHz and $E_{\text{up}} = 323.7$ K between 1.7 and 3.7 km s^{-1} . The positions of the continuum of source A (South East source) and source B (North West source) are indicated with red crosses, while the positions studied in this paper are indicated with red circles. The contour levels start at 5σ with additional steps of 5σ . The circular restoring beam of $0.5''$ size is indicated in grey in the right hand lower corner.

$3 \times 10^{15} \text{ cm}^{-2}$ and $4 \times 10^{15} \text{ cm}^{-2}$ for $T_{\text{ex}} = 300$ and 100 K, respectively. All lines are optically thin. It is expected that for the same assumptions all column densities toward source B are accurate to better than 30%.

Toward source A, most of the lines are blended due to the larger linewidths ($\sim 2.5 \text{ km s}^{-1}$). We can, however, identify 11 unblended lines of CH_3NCO (see Figure 4). A column density of $\sim 6 \times 10^{15} \text{ cm}^{-2}$ and $\sim 9 \times 10^{15} \text{ cm}^{-2}$ (assuming a source size of $0.5''$) is in good agreement with observations for excitation temperatures of 300 and 100 K, respectively, again with an uncertainty of about 30%.

At high excitation temperatures, rotational levels in the first excited bending state ($\nu_b=1$) may be populated as well and predictions for possible transitions are shown in Figure B1. For $T_{\text{ex}} = 300$ K, some faint lines can indeed be tentatively attributed to CH_3NCO $\nu_b=1$ transitions toward source B. An integrated intensity map of one of the brightest CH_3NCO lines, the $39_{0,39,0}-38_{0,38,0}$ transition at 336339.9 MHz, is presented in Figure 2. Similarly to other complex molecules, the emission is quite compact with a size of ~ 60 AU radius and centred near the two sources, with little difference between them (Coutens et al. 2016, Jørgensen et al. 2016, Lykke et al. 2017). For source B, the emission is somewhat offset due to the continuum becoming optically thick on source.

We also searched for spectral evidence of two isomers of methyl isocyanate - acetonitrile oxide (CH_3CNO , Winnenwiser et al. 1982) and methyl cyanate (CH_3OCN , Sakaizumi et al. 1990, Kolesniková et al. 2016) - but the corresponding transitions were not found in the PILS data set. From this non-detection 3σ upper limits of $5 \times 10^{13} \text{ cm}^{-2}$

and $5 \times 10^{14} \text{ cm}^{-2}$ are derived, respectively (assuming an excitation temperature of 100 K). These isomers are consequently less abundant than methyl isocyanate by at least two and one orders of magnitude, in agreement with recent findings toward Orion KL (Cernicharo et al. 2016, Kolesniková et al. 2016).

In view of the important role that HNCO and HOCN may play in the formation of CH_3NCO , we also searched for spectral signatures of these precursor species. The analysis of the PILS data for HNCO toward source B is presented in Coutens et al. (2016). The HNCO lines are optically thick, so lines of the isotopologue HN^{13}CO were analysed and the HNCO column density was derived assuming a $^{12}\text{C}/^{13}\text{C}$ ratio of 68 (Milam et al. 2005). To get precise abundance ratios, we re-analysed the HN^{13}CO data using the same dataset (with the circular restoring beam of $0.5''$) and obtained a HNCO column density of $3.7 \times 10^{16} \text{ cm}^{-2}$, which is within the 30% uncertainty range. The resulting HNCO/ CH_3NCO abundance ratio is about 12 toward source B with an uncertainty of less than a factor 2. Within the error margins, this is similar to the value derived in Orion KL (Cernicharo et al. 2016) or Sgr B2(N2) (Belloche et al. 2016, 2017). For source A, the column density of HNCO is estimated to be about $3.4 \times 10^{16} \text{ cm}^{-2}$ ($T_{\text{ex}} = 100$ K). The corresponding HNCO/ CH_3NCO ratio is consequently about 4 toward this component, with a somewhat larger uncertainty of a factor of 3 due to the difficulty of deriving a precise column density for HNCO because of line blending. Overall, we can conclude that the two HNCO/ CH_3NCO ratios are comparable toward the two components of the binary.

Whereas HNCO is readily identified, this is not the case for HOCN. No HOCN lines were detected and a 3σ upper limit of $3 \times 10^{13} \text{ cm}^{-2}$ is derived toward source B. HOCN is consequently less abundant than HNCO and CH_3NCO by at least 3 and 2 orders of magnitude, respectively. Therefore, it is highly unlikely that the gas-phase formation pathway involving HOCN, as proposed by Halfen et al. (2015) in Eq. 1 and 2, contributes significantly to the production of methyl isocyanate in this source.

We have also determined the abundance ratios of CH_3NCO with respect to CH_3OH and CH_3CN for comparison with other sources in which methyl isocyanate has been detected (see Table 1). The column density of CH_3OH was estimated based on $\text{CH}_3^{18}\text{OH}$ by Jørgensen et al. (in prep.) for the one beam offset position toward source B ($\sim 1 \times 10^{19} \text{ cm}^{-2}$) using a $^{16}\text{O}/^{18}\text{O}$ ratio of 560 (Wilson & Rood 1994). With the same assumptions, we estimate a column density of CH_3OH of $\sim 2 \times 10^{19} \text{ cm}^{-2}$ in source A. CH_3CN was analysed by Calcutt et al. (in prep.) toward both source A ($\sim 8 \times 10^{16} \text{ cm}^{-2}$) and source B ($\sim 3 \times 10^{16} \text{ cm}^{-2}$). Similarly to the HNCO/ CH_3NCO ratio, the abundance ratio of $\text{CH}_3\text{CN}/\text{CH}_3\text{NCO}$ is comparable to the values found in Orion KL (Cernicharo et al. 2016) and lower than toward Sgr B2(N1), but again comparable to Sgr B2(N2) (Belloche et al. 2016, 2017). Methanol is, however, clearly more abundant than methyl isocyanate toward IRAS16293 than toward Orion KL and Sgr B2.

The HNCO/ CH_3NCO and $\text{CH}_3\text{CN}/\text{CH}_3\text{NCO}$ abundance ratios derived in IRAS16293 are much larger than the lower limits found in comet 67P (>0.2). A proper com-

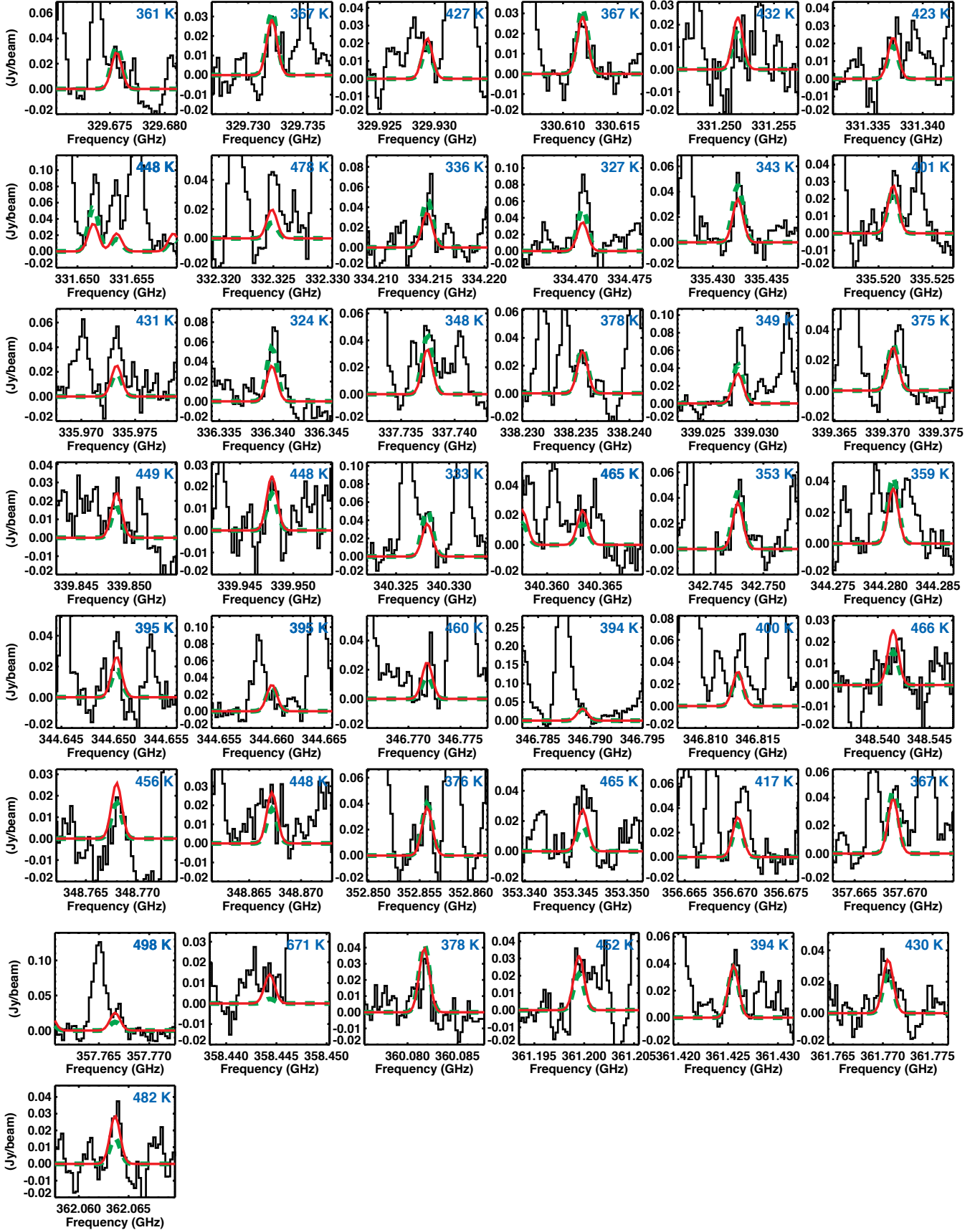


Figure 3. Black: detected unblended lines of CH_3NCO toward source B. Red solid: best-fit model for $T_{\text{ex}} = 300$ K. Green dashed: best-fit model for $T_{\text{ex}} = 100$ K. The E_{up} values of the lines are indicated in blue in the right upper part of each panel.

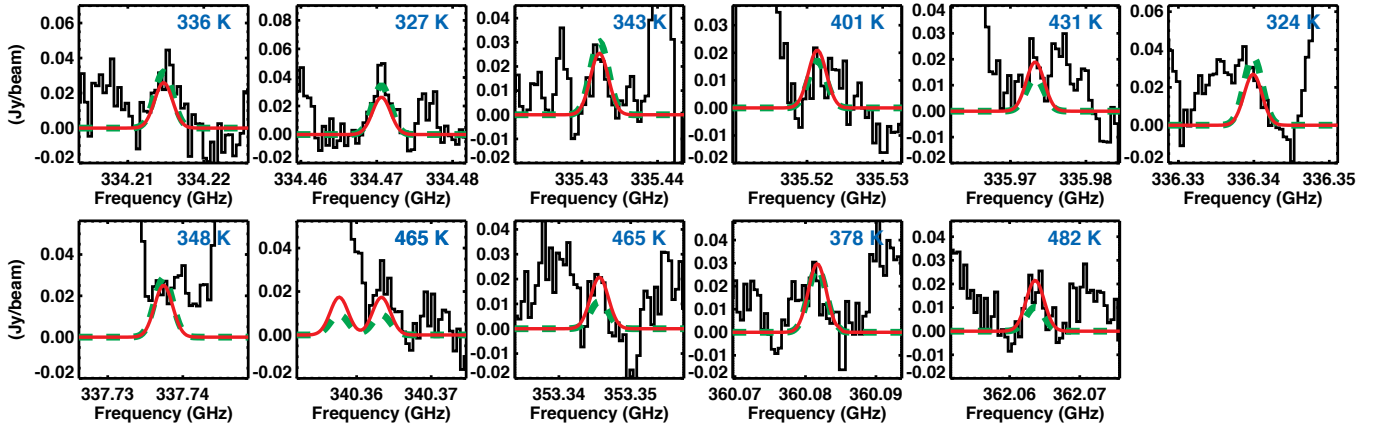


Figure 4. Black: detected unblended lines of CH₃NCO toward source A. Red solid: best-fit model for $T_{\text{ex}} = 300$ K. Green dashed: best-fit model for $T_{\text{ex}} = 100$ K. The E_{up} values of the lines are indicated in blue in the right upper part of each panel.

Table 1. Comparison of molecular abundance ratios toward different sources

Source	HNCO/CH ₃ NCO	CH ₃ CN/CH ₃ NCO	CH ₃ OH/CH ₃ NCO	Reference
IRAS16293 B	12	8	3333	This work
IRAS16293 A	4	9	3200	
Orion KL A	15	10	400	Cernicharo et al. (2016)
Orion KL B	15	25	225	
Sgr B2(N1)	40	50	40	
Sgr B2(N2)	9	10	182	Belloche et al. (2016, 2017)
Comet 67P/C-G	>0.2	>0.2	–	Goesmann et al. (2015) Altwegg et al. in prep.

parison awaits the new results derived with the ROSINA instrument (Altwegg et al. in prep.).

3 LABORATORY EXPERIMENTS

Grain surface formation routes of complex molecules have been investigated experimentally for many years using cryogenic set-ups to mimic specific chemical processes under fully controlled laboratory conditions (see Herbst & van Dishoeck 2009, Linnartz et al. 2015, for reviews). In the present work the formation of methyl isocyanate is investigated by energetically processing solid-state CH₄:HNCO mixtures with vacuum-UV radiation. VUV irradiated methane ice is known to efficiently yield methyl radicals (Bossa et al. 2015), and these radicals are expected to further react through surface diffusion with HNCO to form CH₃NCO, as proposed by Goesmann et al. (2015) and Cernicharo et al. (2016), reaction 4. The Cryogenic Photoproduct Analysis Device 2 (CryoPAD2) in the Sackler Laboratory for Astrophysics is used to perform the measurements to investigate the role of methylation of HNCO in methyl isocyanate formation. A short description of this set-up, experimental procedure and reference data is given below.

Note that in these experiments CH₄ is used as the parent of CH₃ but in interstellar space methyl radicals active in

the ice may also originate from CH₃OH dissociation (Öberg et al. 2009) or from direct CH₃ accretion from the gas-phase. The main aim of this section is to investigate whether or not the proposed solid-state reaction as shown in reaction 4 can proceed at temperatures typical for dark cloud conditions in the ISM, i.e., independent of the exact precursor species.

3.1 Set-up and method

CryoPAD2 consists of a central chamber operating under ultra-high vacuum conditions ($P \approx 10^{-10}$ mbar), to which various instruments are attached. On the top a cryostat is mounted which cools a gold-coated reflective surface down to 12 K. The temperature of this surface can be controlled through resistive heating and varied between 12 and 300 K, with an absolute temperature accuracy better than 1 K. Pure and mixed gases are prepared separately in a gas-mixing system which is attached to a high-precision leak valve to the main chamber and which doses the gases onto the cooled surface. The deposited ice samples are irradiated with VUV radiation from a Microwave Discharge Hydrogen-flow Lamp (MDHL), which emits radiation peaking at 121 nm and a region between 140 to 160 nm, corresponding to photon energies in the range of 7.5 to 10.2 eV (Chen et al. 2014, Ligterink et al. 2015). The total lamp flux is $(1.1 \pm 0.1) \times 10^{14}$ photons s⁻¹ that is determined using a cal-

ibrated NIST diode. CryoPAD2 possesses two diagnostic tools to monitor the VUV induced processes in the ice. The beam of a Fourier-Transform IR Spectrometer (FTIRS) is directed under grazing incidence onto the gold-coated surface, in order to perform Reflection Absorption IR Spectroscopy (RAIRS). The sample preparation and chemical changes under influence of VUV radiation are monitored by recording vibrational fingerprint spectra of molecules on the surface. To decrease the perturbing role of atmospheric CO_2 and H_2O , the system is purged with filtered dry air. The second instrument is a highly sensitive Quadrupole Mass Spectrometer (QMS), with an ionization element at 70 eV, which is able to trace gas-phase molecules in the chamber that are released from the ice surface upon linear heating during a Temperature Programmed Desorption (TPD) experiment. TPD is a very useful method that allows to identify desorbing species through their specific desorption temperature and mass fragmentation pattern. Unambiguous identifications become possible through the use of isotopologues and searching for the corresponding mass shifts of specific fragments. Obviously, TPD comes with the destruction of the ice.

In the experiments CH_4 (Linde Gas, 99.995% purity), $^{13}\text{CH}_4$ (Sigma-Aldrich, 99% purity) and HNCO are used. HNCO is produced from cyranic acid (Sigma-Aldrich, 98% purity), the solid trimer of HNCO, following the protocol described in van Broekhuizen et al. (2004). Impurities of the HNCO production process are removed by freeze-thaw cycles to obtain a HNCO purity of >99%.

For the experiments samples of pure HNCO and methane, and mixtures of $^{13}\text{CH}_4/\text{CH}_4:\text{HNCO}$ at 5:1 ratio are prepared. This ratio is within a factor of 2 of that observed for interstellar ices (Öberg et al. 2011) but is particularly chosen to create a large abundance of CH_3 radicals to test whether reaction 4 proceeds or not. Homogeneously mixed ices are grown on the surface at 20 K and irradiated with a total fluence of $\sim 3.3 \times 10^{17}$ photons. During irradiation of the sample, IR spectra are continuously recorded at 1 cm^{-1} resolution. After the irradiation TPD is started, while still recording IR spectra.

The strongest vibrational features of solid methyl isocyanate are found between 2320 and 2280 cm^{-1} ($\sim 4.34 \mu\text{m}$) for the $-\text{N}=\text{C}=\text{O}$ antisymmetric stretching vibration and overtone $2\nu_7$ CH_3 rocking mode. Sullivan et al. (1994) lists these at 2320, 2280, 2270, 2240 and 2228 cm^{-1} , with 2280 cm^{-1} being the strongest band. Zhou & Durig (2009) positioned all bands around 2300 cm^{-1} and Reva et al. (2010) puts the band for methyl isocyanate in a nitrogen matrix at 2334.7, 2307.9, 2288.9, 2265.2 and 2259.7 cm^{-1} , finding the strongest transition at 2288.9 cm^{-1} . The region around these bands is used to monitor CH_3NCO formation in the ice. We focus on the region between 2400 – 2100 cm^{-1} . From previous experiments it is known that CO_2 (2340 cm^{-1}), OCN^- (2165 cm^{-1}) and CO (2135 cm^{-1}) are produced from HNCO (2265 cm^{-1}) upon irradiation and these photo-products also have spectral features in the region characterizing methyl isocyanate (Raunier et al. 2004, van Broekhuizen et al. 2004). Other known products of HNCO irradiation are formamide, urea and formaldehyde, which do not have any interfering IR features in the region of interest. Energetic processing of methane does not yield products that show up in the region of interest. Products that are formed from

methane are mainly ethane and to a lesser extent ethene and ethyn (Bennett et al. 2006, Bossa et al. 2015). These species are seen in our spectra at their appropriate frequencies. The abundant formation of ethane demonstrates that CH_3 is indeed produced in the experiments, since ethane is the direct product of methyl-radical recombination.

In order to identify methyl isocyanate in the gas-phase using TPD, the mass fragmentation pattern available from the NIST database² is used. The fragmentation pattern at 70 eV comprises unique peaks at $m/z = 57$ and 56 (hereafter also called the primary and secondary mass peak), which have a m/z 57:56 ratio of 5:2 and these will be used as main TPD mass tracers. Known products of pure HNCO and methane irradiation does not have a mass fragmentation peak at $m/z = 57$ (see also Appendix C).

It is important to mention that methyl isocyanate is severely toxic and specialised laboratories and equipment are needed to work with this substance. This complicates the extensive use of this species and for this reason additional experiments, starting from the pure precursor, have not been performed.

3.2 Results – IR Spectra

Figure 5 presents the IR spectra recorded during the first 1×10^{17} photons irradiating the $^{12/13}\text{CH}_4:\text{HNCO}$ samples. All spectra are normalized to the HNCO peak. Three known spectroscopic features of CO_2 , OCN^- and CO (blue) show up during irradiation. Also, two new features become visible around 2300 cm^{-1} (red), which do not show up while processing samples of pure HNCO or CH_4 . Also a clear red shift of about 10 cm^{-1} of the two features is seen between the sample of $^{12}\text{CH}_4$ and $^{13}\text{CH}_4$, moving transitions at 2322 and 2303 cm^{-1} to 2313 and 2294 cm^{-1} . These spectroscopic features are therefore the result of a product formed in the reaction between methane and isocyanic acid, and, since they are found close to known CH_3NCO features (given by Sullivan et al. (1994) and Reva et al. (2010)), are plausibly identified with methyl isocyanate. Another feature is seen in the wing of the HNCO peak around 2235 cm^{-1} , which does not clearly shift with methane isotopologues. The origin of this band is unclear.

Bandstrength values for methyl isocyanate are not available from the literature, however, a rough indication of the amount of formed methyl isocyanate versus lost HNCO can be given by making the assumption that the bandstrength of the NCO antisymmetric stretch vibration of methyl isocyanate equals that of the corresponding vibration of HNCO. To obtain the ratio, the integrated area of the 2303 cm^{-1} feature is divided by the integrated loss area of the 2265 cm^{-1} HNCO band for a number of spectra. A ratio of $N(\text{HNCO})/N(\text{CH}_3\text{NCO}) = 100 - 200$ is found, which is about an order of magnitude higher than the ratio observed toward IRAS16293. It should be noted, that it is not *a priori* clear whether solid state laboratory and gas phase

² NIST Mass Spec Data Center, S.E. Stein, director, "Mass Spectra" in NIST Chemistry WebBook, NIST Standard Reference Database Number 69, Eds. P.J. Lindstrom and W.G. Mallard, National Institute of Standards and Technology, Gaithersburg MD, 20899, <http://webbook.nist.gov>.

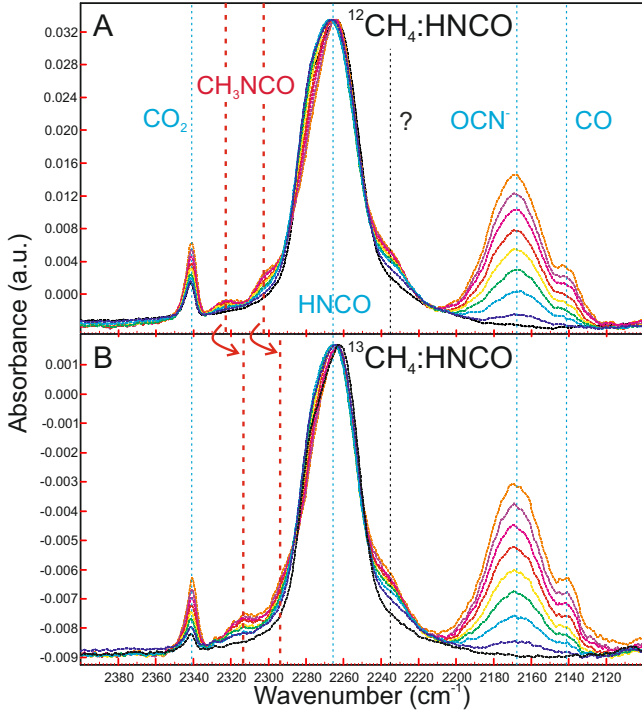


Figure 5. IR spectra taken over time for the $^{12}\text{CH}_4\text{:HNCO}$ (A) and $^{13}\text{CH}_4\text{:HNCO}$ (B) mixture. HNCO and the products CO_2 , OCN^- and CO are listed (blue). Spectroscopic features that coincide with CH_3NCO are found at the red lines and show a clear shift with the $^{13}\text{CH}_4$ isotopologue. One unidentified peak is found in the right wing of the HNCO peak, indicated by a question mark.

astronomical abundances can be directly compared (see e.g. Chuang et al. 2017). It is likely that ongoing gas phase reactions, for one or the other species, change the solid state to gas phase ratios. It is also possible that non-linear RAIRS effects can offset the column density (Teolis et al. 2007) or that the CH_3NCO bandstrength is significantly different. In a same manner CH_3NCO photodestruction may affect the overall abundances.

3.3 Results – temperature programmed desorption

The desorption temperature of methyl isocyanate has not been reported in the literature, but the TPD traces of our experiments on UV processed $\text{CH}_4\text{:HNCO}$ ices do show the combined release of the primary and secondary masses of methyl isocyanate $m=57$ and 56 at 145 K (Figure 6). The verification experiment of pure HNCO did not show these masses and only $m/z = 56$ (and no 57) was seen after the irradiation of pure methane ice, releasing at 105 K (Figure C1 in the Appendix). At the 145 K desorption peak the primary over secondary mass ratio is around $1\text{--}1.5$, lower than the value of 2.5 suggested by NIST upon 70 eV electron impact ionization. The NIST calibration values are a good indicator of the values to be expected, but are to some extent set-up specific. In such cases, the expected reactant

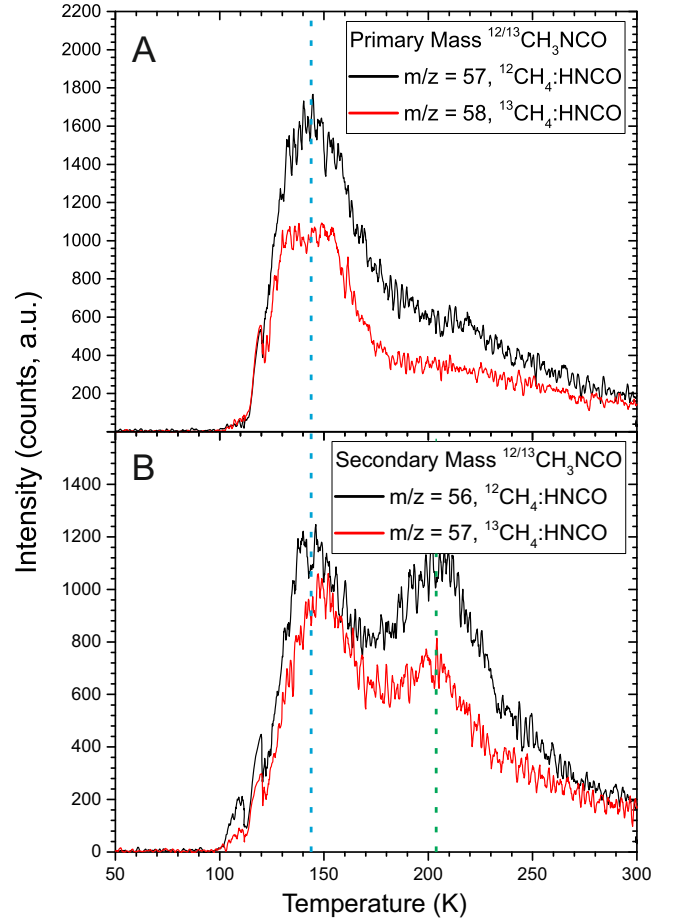


Figure 6. TPD trace of the primary (A) and secondary (B) masses of $^{12/13}\text{CH}_3\text{NCO}$. Both the primary and secondary mass are seen being released around 145 K (blue line). The secondary mass shows another release peak around 205 K (green), which is suspected to also contribute to first release peak, thus altering respectively the m/z $57/56$ and $58/57$ ratio.

can be deposited directly and the fragmentation pattern can be studied adapted to the set-up in use. This is unfortunately not possible here due to the aforementioned toxicity of methyl isocyanate. Moreover, contributions from other reaction products cannot be fully excluded. In Figure 6B a second desorption peak is found around 205 K , which is seen only as a shoulder in panel A. It is unlikely that this peak is associated with methyl isocyanate, instead it shows that the chemical network involves the formation of other species as well. In fact, since the additional product(s) have a mass fragmentation pattern which contributes to the secondary mass of methyl isocyanate, this could explain why the primary/secondary mass ratio does not exactly match with that given by NIST. Additionally, the unidentified product of pure methane irradiation could be trapped in the ice and contribute to other desorption peaks. Therefore the 145 K desorption peak is still consistent with methyl isocyanate. There is also no other candidate molecule with a primary mass of 57 in the NIST database that could plausibly explain the TPD spectra.

CryoPAD2 offers the unique feature to combine IR and

TPD data, i.e., at the moment a specific ice feature starts thermally desorbing, the IR signal starts decreasing and simultaneously the mass signal is increasing. In the measurements presented here, this effect is not as clearly visible as in previous studies (e.g. Öberg et al. (2009)), but this may be partly due to the low final S/N of the feature, making it difficult to link RAIRS and TPD signals. The present data are as accurate as possible, but obviously their interpretation would benefit from focused experiments determining the thermal desorption peak of pure CH_3NCO and the IR band strengths.

4 LINKING OBSERVATIONS AND LABORATORY DATA

Methyl isocyanate is detected in a significant abundance toward both components of IRAS16293 with an excitation temperature of at least 100 K. The two isomers, CH_3CNO and CH_3OCN , are not detected. Abundance ratios with respect to HNCO and CH_3CN are found to be similar to those found for Orion KL and comparable within a factor of a few to Sgr B2, making a common formation pathway in these three sources likely. Moreover, no chemical differentiation between source A and B is found. Gas-phase production of CH_3NCO via HOCN (reaction 1) can be seen as insignificant due to the low upper limit on this molecule, but a gas-phase route involving HNCO cannot be excluded.

The laboratory experiments on the formation of CH_3NCO strongly suggest that a solid-state formation scheme is possible, starting from two astronomically relevant precursor species, CH_4 and HNCO . The proposed reaction $\text{CH}_3 + \text{HNCO} \rightarrow \text{CH}_3\text{NCO} + \text{H}$ indeed seems to proceed, confirming a solid state reaction pathway, but other routes such as hydrogen stripping $\text{CH}_3 + \text{HNCO} \rightarrow \text{CH}_4 + \text{NCO}$ may take place in parallel. Reactions involving the OCN^- anion, which is abundantly formed in these experiments, provide alternative routes. Irradiation of a sample of $\text{OCN}^-:\text{CH}_4$, where OCN^- is formed via the acid-base reaction of HNCO with NH_3 , will be interesting to study in future work since this anion is also a well known component of interstellar ice (Boogert et al. 2015).

Belloche et al. (2017) have incorporated reactions 5 and 6 (see Section 1) into the large gas-grain model MAGICKAL (Garrod 2013) and simulated the cold collapse + warm-up phase of a hot core region. Although tailored to the high-mass source SgrB2(N) many of the model features are also expected to be valid for the case of IRAS16293. In this model, HNCO is formed by reaction of $\text{NH} + \text{CO}$, a route that has been demonstrated experimentally by Fedoseev et al. (2015, 2016). Depending on assumptions about the barriers of the grain surface reactions involved in the formation of HNCO and CH_3NCO , abundance ratios $\text{HNCO}/\text{CH}_3\text{NCO}$ of 3-50 following ice sublimation are readily found in the models, consistent with the IRAS16293 observations.

5 CONCLUSIONS

The main observational and experimental conclusions are listed below:

(i) Methyl isocyanate is detected for the first time toward a low mass protostar, IRAS16293, on solar system scales (emission radius of 60 AU). Column densities of $\sim (3-4) \times 10^{15}$ and $\sim (6-9) \times 10^{15} \text{ cm}^{-2}$ are obtained toward source B and A, respectively, yielding $N(\text{HNCO})/N(\text{CH}_3\text{NCO}) = 12$ and 4, with no significant variation between the two sources.

(ii) The abundance ratios of CH_3NCO relative to the N-bearing species HNCO and CH_3CN are similar to those found toward Orion KL and deviate by at most an order of magnitude from Sgr B2.

(iii) The isomers of methyl isocyanate, CH_3CNO and CH_3OCN , are not detected. These species are less abundant than CH_3NCO by at least a factor 100 and 10, respectively.

(iv) HOCN is not found down to $< 3 \times 10^{13} \text{ cm}^{-2}$, giving $N(\text{HOCN})/N(\text{HNCO}) > 1000$, which makes this an insignificant gas-phase precursor to methyl isocyanate in IRAS16293.

(v) Laboratory experiments strongly suggest that it is possible to form CH_3NCO on an icy surface, irradiating CH_4 and HNCO as astronomically relevant precursor species with VUV light, generating methyl radicals as reactive intermediates to form methyl isocyanate.

The detection of CH_3NCO adds to the growing list of complex molecules known to be present around solar mass protostars, showing that the ingredients for prebiotic molecules are abundant. Future deeper searches for even more complex molecules relevant for the origin of life are warranted.

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REFERENCES

Baryshev A. M., et al., 2015, *A&A*, **577**, A129

- Belloche A., Müller H. S. P., Menten K. M., Schilke P., Comito C., 2013, *A&A*, **559**, A47
- Belloche A., Garrod R. T., Müller H. S. P., Menten K. M., 2014, *Science*, **345**, 1584
- Belloche A., Müller H. S. P., Garrod R. T., Menten K. M., 2016, *A&A*, **587**, A91
- Belloche A., et al., 2017, accepted in *A&A*
- Bennett C. J., Jamieson C. S., Osamura Y., Kaiser R. I., 2006, *ApJ*, **653**, 792
- Bisschop S. E., Jørgensen J. K., van Dishoeck E. F., de Wachter E. B. M., 2007, *A&A*, **465**, 913
- Bisschop S. E., Jørgensen J. K., Bourke T. L., Bottinelli S., van Dishoeck E. F., 2008, *A&A*, **488**, 959
- Blake G. A., Sutton E. C., Masson C. R., Phillips T. G., 1987, *ApJ*, **315**, 621
- Boogert A. C. A., Gerakines P. A., Whittet D. C. B., 2015, *ARA&A*, **53**, 541
- Bossa J.-B., Paardekooper D. M., Isokoski K., Linnartz H., 2015, *Physical Chemistry Chemical Physics (Incorporating Faraday Transactions)*, **17**, 17346
- Bottinelli S., et al., 2004, *ApJ*, **617**, L69
- Cazaux S., Tielens A. G. G. M., Ceccarelli C., Castets A., Wakeham V., Caux E., Parise B., Teyssier D., 2003, *ApJ*, **593**, L51
- Cernicharo J., et al., 2016, *A&A*, **587**, L4
- Chen Y.-J., Chuang K.-J., Muñoz Caro G. M., Nuevo M., Chu C.-C., Yih T.-S., Ip W.-H., Wu C.-Y. R., 2014, *ApJ*, **781**, 15
- Coutens A., et al., 2016, *A&A*, **590**, L6
- Crockett N. R., et al., 2014, *ApJ*, **787**, 112
- Fedoseev G., Ioppolo S., Zhao D., Lamberts T., Linnartz H., 2015, *MNRAS*, **446**, 439
- Fedoseev G., Chuang K.-J., van Dishoeck E. F., Ioppolo S., Linnartz H., 2016, *MNRAS*, **460**, 4297
- Feuchtgruber H., Helmich F. P., van Dishoeck E. F., Wright C. M., 2000, *ApJ*, **535**, L111
- Garrod R. T., 2013, *ApJ*, **778**, 158
- Goesmann F., et al., 2015, *Science*, **349**
- Halfen D. T., Ilyushin V. V., Ziurys L. M., 2015, *ApJ*, **812**, L5
- Henderson B. L., Gudipati M. S., 2015, *ApJ*, **800**, 66
- Herbst E., van Dishoeck E. F., 2009, *ARA&A*, **47**, 427
- Jaber A. A., Ceccarelli C., Kahane C., Caux E., 2014, *ApJ*, **791**, 29
- Jørgensen J. K., Bourke T. L., Nguyen Luong Q., Takakuwa S., 2011, *A&A*, **534**, A100
- Jørgensen J. K., Favre C., Bisschop S. E., Bourke T. L., van Dishoeck E. F., Schmalzl M., 2012, *ApJ*, **757**, L4
- Jørgensen J. K., et al., 2016, *A&A*, **595**, A117
- Kolesniková L., Alonso J. L., Bermúdez C., Alonso E. R., Tercero B., Cernicharo J., Guillemin J.-C., 2016, *A&A*, **591**, A75
- Koput J., 1986, *Journal of Molecular Spectroscopy*, **115**, 131
- Kuan Y.-J., et al., 2004, *ApJ*, **616**, L27
- Ligterink N. F. W., Paardekooper D. M., Chuang K.-J., Both M. L., Cruz-Diaz G. A., van Helden J. H., Linnartz H., 2015, *A&A*, **584**, A56
- Linnartz H., Ioppolo S., Fedoseev G., 2015, *International Reviews in Physical Chemistry*, **34**, 205
- López-Sepulcre A., et al., 2015, *MNRAS*, **449**, 2438
- Lykke J. M., et al., 2017, *A&A*, **597**, A53
- Martín-Doménech R., Rivilla V., Jiménez-Serra I., Quenard D., Testi L., Martín-Pintado J., 2017, preprint, ([arXiv:1701.04376](https://arxiv.org/abs/1701.04376))
- Milam S. N., Savage C., Brewster M. A., Ziurys L. M., Wyckoff S., 2005, *ApJ*, **634**, 1126
- Müller H. S. P., Thorwirth S., Roth D. A., Winnewisser G., 2001, *A&A*, **370**, L49
- Müller H. S. P., Schlöder F., Stutzki J., Winnewisser G., 2005, *Journal of Molecular Structure*, **742**, 215
- Neill J. L., et al., 2014, *ApJ*, **789**, 8
- Nummelin A., Bergman P., Hjalmarsen Å., Friberg P., Irvine W. M., Millar T. J., Ohishi M., Saito S., 2000, *ApJS*, **128**, 213
- Öberg K. I., Garrod R. T., van Dishoeck E. F., Linnartz H., 2009, *A&A*, **504**, 891
- Öberg K. I., Boogert A. C. A., Pontoppidan K. M., van den Broek S., van Dishoeck E. F., Bottinelli S., Blake G. A., Evans II N. J., 2011, *ApJ*, **740**, 109
- Parise B., Simon T., Caux E., Dartois E., Ceccarelli C., Rayner J., Tielens A. G. G. M., 2003, *A&A*, **410**, 897
- Pineda J. E., et al., 2012, *A&A*, **544**, L7
- Raunier S., Chiavassa T., Duvernay F., Borget F., Aycard J. P., Dartois E., d'Hendecourt L., 2004, *A&A*, **416**, 165
- Reva I., Lapinski L., Fausto R., 2010, *J. Mol. Struct.*, **976**, 333
- Ruzi M., Anderson D., 2012, *J. Chem. Phys.*, **137**, 194313
- Sakaizumi T., Mure H., Ohashi O., Yamaguchi I., 1990, *Journal of Molecular Spectroscopy*, **140**, 62
- Sullivan J., Heusel H., Zunic W., Durig J., 1994, *Spectrochimica Acta*, **50**
- Teolis B. D., Loeffler M. J., Raut U., Famá M., Baragiola R. A., 2007, *Icarus*, **190**, 274
- Tercero B., Kleiner I., Cernicharo J., Nguyen H. V. L., López A., Muñoz Caro G. M., 2013, *ApJ*, **770**, L13
- Wilson T. L., Rood R., 1994, *ARA&A*, **32**, 191
- Winnewisser M., Pearson E. F., Galica J., Winnewisser B. P., 1982, *Journal of Molecular Spectroscopy*, **91**, 255
- Zhou S., Durig J., 2009, *JMS*, **924**
- van Broekhuizen F. A., Keane J. V., Schutte W. A., 2004, *A&A*, **415**, 425
- van Broekhuizen F. A., Pontoppidan K. M., Fraser H. J., van Dishoeck E. F., 2005, *A&A*, **441**, 249
- van Dishoeck E. F., Blake G. A., Jansen D. J., Groesbeck T. D., 1995, *ApJ*, **447**, 760

APPENDIX A: LINE LIST

Table A1 contains the full list of detected CH_3NCO , $v_B=0$ transitions toward IRAS16293. A total of 43 lines are observed.

APPENDIX B: $\text{CH}_3\text{NCO } v_B=1$

Figure B1 shows the tentative identification of the first excited state of methyl isocyanate.

APPENDIX C: $M/Z = 56, 57$ AND 58 UPON VUV IRRADIATION OF PURE CH_4 AND HNCO SAMPLES

Several publications address the energetic processing of pure HNCO and CH_4 ice and much is known about the chemistry that can be induced in these species (Raunier et al. 2004, Bennett et al. 2006, Bossa et al. 2015). Of the known chemical products, none of them contributes to the primary and secondary masses of methyl isocyanate. To verify this, pure methane and isocyanic acid have been VUV irradiated. The HNCO sample (Figure C1A) does not show any significant release of these masses. However, after irradiation of methane a substantial amount of $m/z = 56$, coinciding with a lesser amount of $m/z = 58$, is seen releasing at a desorption temperature of 105 K. Some $m/z = 57$ is seen to be released as well, although there is no clear peak found in this case.

Table A1. Unblended lines of CH₃NCO detected toward source B.

Species	Transition	Frequency (MHz)	E_{up} (K)	A_{ij} (s ⁻¹)	g_{up}
CH ₃ NCO, $\nu_b=0$	(38 0 0 2 – 37 0 0 2)	329675.6	361.3	1.71×10^{-3}	77
CH ₃ NCO, $\nu_b=0$	(38 1 0 2 – 37 1 0 2)	329732.1	367.3	1.71×10^{-3}	77
CH ₃ NCO, $\nu_b=0$	(38 -1 0 -3 – 37 -1 0 -3)	329929.3	427.1	1.67×10^{-3}	77
CH ₃ NCO, $\nu_b=0$	(38 -1 0 2 – 37 -1 0 2)	330611.8	367.3	1.71×10^{-3}	77
CH ₃ NCO, $\nu_b=0$	(38 1 0 3 – 37 1 0 3)	331251.7	431.6	1.74×10^{-3}	77
CH ₃ NCO, $\nu_b=0$	(38 0 0 -3 – 37 0 0 -3)	331337.3	423.4	1.70×10^{-3}	77
CH ₃ NCO, $\nu_b=0$	(38 2 0 3 – 37 2 0 3)	331653.6	447.5	1.72×10^{-3}	77
CH ₃ NCO, $\nu_b=0$	(38 3 0 -3 – 37 3 0 -3)	332324.9	477.7	1.71×10^{-3}	77
CH ₃ NCO, $\nu_b=0$	(39 0 0 1 – 38 0 0 1)	334214.7	336.5	1.82×10^{-3}	79
CH ₃ NCO, $\nu_b=0$	(39 1 39 0 – 38 1 38 0)	334470.6	327.3	1.78×10^{-3}	79
CH ₃ NCO, $\nu_b=0$	(39 -1 0 1 – 38 -1 0 1)	335432.3	342.5	1.83×10^{-3}	79
CH ₃ NCO, $\nu_b=0$	(39 2 0 2 – 38 2 0 2)	335521.3	401.3	1.83×10^{-3}	79
CH ₃ NCO, $\nu_b=0$	(39 3 0 2 – 38 3 0 2)	335973.3	431.1	1.82×10^{-3}	79
CH ₃ NCO, $\nu_b=0$	(39 0 39 0 – 38 0 38 0)	336339.9	323.7	1.82×10^{-3}	79
CH ₃ NCO, $\nu_b=0$	(39 2 38 0 – 38 2 37 0)	337737.4	348.4	1.83×10^{-3}	79
CH ₃ NCO, $\nu_b=0$	(39 0 0 2 – 38 0 0 2)	338235.6	377.6	1.85×10^{-3}	79
CH ₃ NCO, $\nu_b=0$	(39 2 37 0 – 38 2 36 0)	339028.2	349.0	1.85×10^{-3}	79
CH ₃ NCO, $\nu_b=0$	(38 3 0 1 – 37 3 0 1)	339370.5	374.5	1.75×10^{-3}	77
CH ₃ NCO, $\nu_b=0$	(39 1 0 -3 – 38 1 0 -3)	339848.8	448.8	1.88×10^{-3}	79
CH ₃ NCO, $\nu_b=0$	(39 1 0 3 – 38 1 0 3)	339948.0	448.0	1.88×10^{-3}	79
CH ₃ NCO, $\nu_b=0$	(39 1 38 0 – 38 1 37 0)	340327.9	333.1	1.88×10^{-3}	79
CH ₃ NCO, $\nu_b=0$	(39 2 0 -3 – 38 2 0 -3)	340363.3	464.6	1.87×10^{-3}	79
CH ₃ NCO, $\nu_b=0$	(40 0 0 1 – 39 0 0 1)	342747.6	353.1	1.97×10^{-3}	81
CH ₃ NCO, $\nu_b=0$	(40 1 0 1 – 39 1 0 1)	344280.7	359.2	1.98×10^{-3}	81
CH ₃ NCO, $\nu_b=0$	(40 3 0 2 – 39 3 0 2)	344650.3	447.7	1.97×10^{-3}	81
CH ₃ NCO, $\nu_b=0$	(40 3 38 0 – 39 3 37 0)	344660.0	395.0	1.97×10^{-3}	81
CH ₃ NCO, $\nu_b=0$	(40 -1 0 -3 – 39 -1 0 -3)	346771.8	459.6	1.95×10^{-3}	81
CH ₃ NCO, $\nu_b=0$	(40 0 0 2 – 39 0 0 2)	346789.2	394.2	1.99×10^{-3}	81
CH ₃ NCO, $\nu_b=0$	(40 1 0 2 – 39 1 0 2)	346813.1	400.2	1.99×10^{-3}	81
CH ₃ NCO, $\nu_b=0$	(40 1 0 -3 – 39 1 0 -3)	348541.5	465.6	2.03×10^{-3}	81
CH ₃ NCO, $\nu_b=0$	(40 0 0 -3 – 39 0 0 -3)	348768.0	456.2	1.99×10^{-3}	81
CH ₃ NCO, $\nu_b=0$	(40 -3 0 2 – 39 -3 0 2)	348867.2	447.9	1.99×10^{-3}	81
CH ₃ NCO, $\nu_b=0$	(41 1 0 1 – 40 1 0 1)	352855.7	376.2	2.13×10^{-3}	83
CH ₃ NCO, $\nu_b=0$	(41 3 0 2 – 40 3 0 2)	353345.6	464.8	2.12×10^{-3}	83
CH ₃ NCO, $\nu_b=0$	(41 -1 0 2 – 40 -1 0 2)	356670.2	417.3	2.15×10^{-3}	83
CH ₃ NCO, $\nu_b=0$	(41 1 40 0 – 40 1 39 0)	357668.8	367.0	2.18×10^{-3}	83
CH ₃ NCO, $\nu_b=0$	(41 2 0 -3 – 40 2 0 -3)	357766.6	498.4	2.17×10^{-3}	83
CH ₃ NCO, $\nu_b=0$	(41 0 0 5 – 40 0 0 5)	358444.3	671.4	2.16×10^{-3}	83
CH ₃ NCO, $\nu_b=0$	(42 1 42 0 – 41 1 41 0)	360081.7	378.0	2.23×10^{-3}	85
CH ₃ NCO, $\nu_b=0$	(42 2 0 2 – 41 2 0 2)	361199.5	452.4	2.28×10^{-3}	85
CH ₃ NCO, $\nu_b=0$	(42 1 0 1 – 41 1 0 1)	361425.5	393.7	2.29×10^{-3}	85
CH ₃ NCO, $\nu_b=0$	(42 3 40 0 – 41 3 39 0)	361770.5	429.5	2.28×10^{-3}	85
CH ₃ NCO, $\nu_b=0$	(42 3 0 2 – 41 3 0 2)	362063.6	482.2	2.28×10^{-3}	85

The resulting contribution of $m/z = 56$ can interfere with the $m/z = 57/56$ amu mass ratio of methyl isocyanate.

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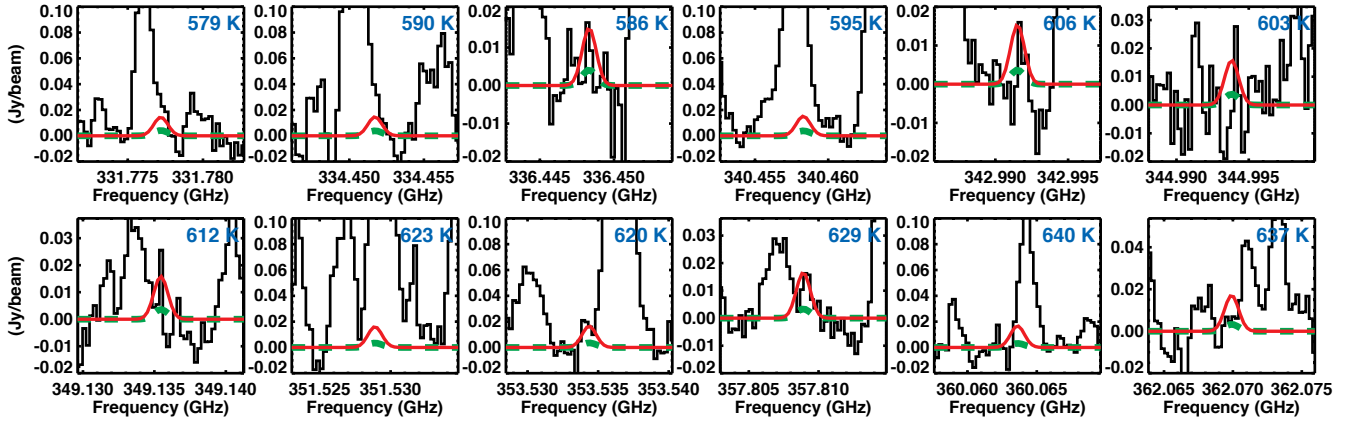


Figure B1. Predictions for the $\text{CH}_3\text{NCO } \nu_b=1$ lines toward source B (red solid: best-fit model for $T_{\text{ex}} = 300$ K, green dashed: best-fit model for $T_{\text{ex}} = 100$ K).

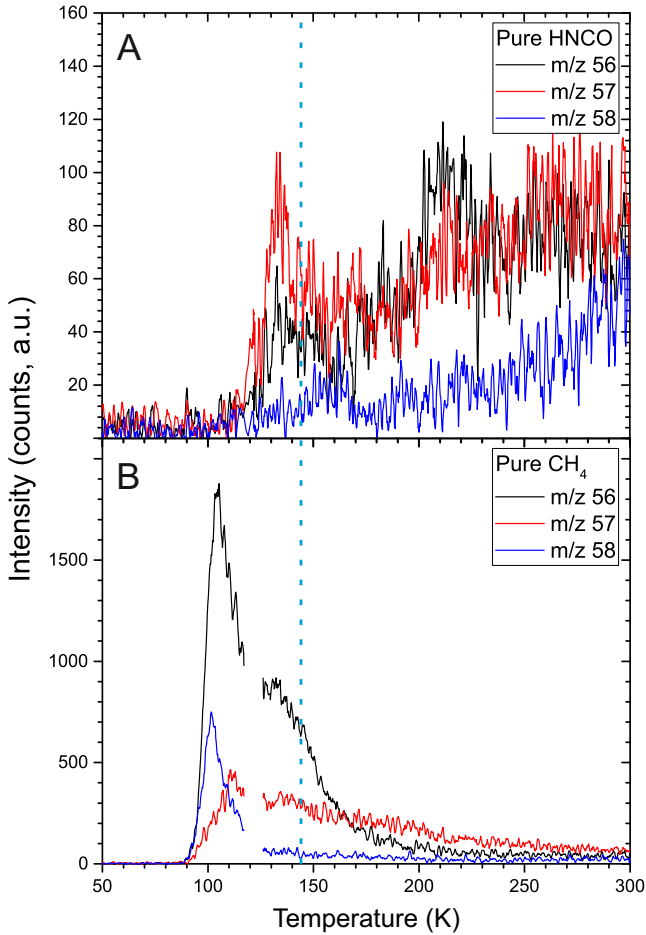


Figure C1. TPD trace of $m/z = 57$ (primary mass CH_3NCO), $m/z = 56$ (secondary mass CH_3NCO) and $m/z = 58$ after the irradiation of pure HNCO (A) and pure CH_4 (B). The desorption peak found in the CH_4 :HNCO mixtures is indicated by the dashed line at 145 K.